

Table 2
HyGAL Source Parameters

#	Source Designation	R.A. (hh:mm:ss)	Decl. (dd:mm:ss)	Gal. Long. (deg)	Gal. Lat. (deg)	v_{LSR} (km s ⁻¹)	d [Ref] (kpc)	R_{GAL} (kpc)
a	HGAL284.015–00.86	10:20:16.1	–58:03:55.0	284.016	–0.857	9.0	5.7 [1]	9.0
b	HGAL285.26–00.05	10:31:29.5	–58:02:19.5	285.263	–0.051	3.4	4.3 [2]	8.2
c	G291.579–00.431	11:15:05.7	–61:09:40.8	291.579	–0.431	13.6	8.0 [3]	9.3
d	IRAS 12326-6245	12:35:35.9	–63:02:29.0	301.138	–0.225	–39.3	4.6 [4]	7.2
e	G327.3–00.60	15:53:05.0	–54:35:24.0	327.304	–0.551	–46.9	3.1 [5]	6.2
f	G328.307+0.423	15:54:07.2	–53:11:40.0	328.309	+0.429	–93.6	5.8 [6]	4.6
g	IRAS 16060–5146	16:09:52.4	–51:54:58.5	330.953	–0.182	–91.2	5.3 [7]	4.5
h	IRAS 16164–5046	16:20:11.9	–50:53:17.0	332.827	–0.551	–57.3	3.6 [8]	5.4
i	IRAS 16352–4721	16:38:50.6	–47:28:04.0	337.404	–0.403	–41.4	12.3 [4]	5.1
j	IRAS 16547–4247	16:58:17.2	–42:52:08.9	343.126	–0.063	–30.6	2.7 [8]	5.8
k	NGC 6334 I	17:20:53.4	–35:47:01.5	351.417	+0.645	–7.4	1.3 [9]	7.0
l	G357.558–00.321	17:40:57.2	–31:10:59.3	357.557	–0.321	5.3	9.0–11.8 [10]	1.0–3.6
m	HGAL0.55–0.85	17:50:14.5	–28:54:30.7	0.546	–0.851	16.7	7.7–9.2 [11]	0.4–1.0
n	G09.62+0.19	18:06:14.9	–20:31:37.0	9.620	+0.194	4.3	5.2 [12]	3.3
o	G10.47 + 0.03	18:08:38.4	–19:51:52.0	10.472	+0.026	67.6	8.6 [13]	1.6
p	G19.61–0.23	18:27:38.0	–11:56:39.5	19.608	–0.234	40.8	12.6 [14]	4.7
q	G29.96–0.02	18:46:03.7	–02:39:21.2	29.954	–0.016	97.2	6.7 [6]	4.5
r	G31.41+0.31	18:47:34.1	–01:12:49.0	31.411	+0.307	98.2	4.9 [15]	5.0
s	W43 MM1	18:47:47.0	–01:54:28.0	30.817	–0.057	97.8	5.5 [15]	5.0
t	G32.80+0.19	18:50:30.6	–00:02:00.0	32.796	+0.191	14.6	13.0 [16]	7.4
u	G45.07+0.13	19:13:22.0	+10:50:54.0	45.071	+0.133	59.2	4.3 [17]	6.2
v	DR21	20:39:01.6	+42:19:37.9	81.681	0.537	–4.0	1.5 [18]	7.4
w	NGC 7538 IRS1	23:13:45.3	+61:28:11.7	111.542	0.777	–59.0	2.6 [19]	9.8
x	W3 IRS5	02:25:40.5	+62:05:51.0	133.715	1.215	–39.0	2.3 [20]	9.9
y	W3(OH)	02:27:04.1	+61:52:22.1	133.948	1.064	–48.0	2.0 [20, 21]	9.6

Notes. Those entries highlighted in bold text refer to the sources discussed in this work the remaining sources have either only recently been observed or are yet to be observed.

References. For the heliocentric distances: [1] Urquhart et al. (2014b); [2] Caswell & Vaile (1995); [3] Lee et al. (2012); [4] Green & McClure-Griffiths (2011), [5] Wienen et al. (2015); [5] Urquhart et al. (2018); [7] Moisés et al. (2011); [8] Giannetti et al. (2014); [9] Immer et al. (2013); [10] Frail et al. (1996) (Lacking distance estimates, for G357.558–00.321 we adopt a distance equivalent to the nearby supernova remnant G357.7–0.1 (MSH 17–39)); [11] Walsh et al. (1998); [12] Sanna et al. (2009); [13] Sanna et al. (2014); [14] Urquhart et al. (2014a); [15] Zhang et al. (2014) (For G31.41+0.31, following Winkel et al. (2017) we adopt a distance which is the average between two sources with similar LSR velocities, G31.28+0.06 and G31.58+0.07 for which accurate distances were measured by Zhang et al. (2014) and consistent with Reid et al. (2017).); [16] Kolpak et al. (2003); [17] Stead & Hoare (2010); [18] Rygl et al. (2012); [19] Moscadelli et al. (2009); [20] Navarete et al. (2019); [21] Hachisuka et al. (2006). These distances were also cross-checked with those recently calculated by Mège et al. (2021).